

ring (which they do in a very leisurely manner), leaving the apparent displacement of the planet *on the sun's disc* to subsequent calculation, on a comparison of reports from all the points of observation selected. One-fourth of the advantage arising from its proximity, it is true, is lost, by the sun itself sharing to that extent in the displacement of the planet; but enough remains to give this a superiority over every other method of measuring the sun's distance.

(18.) Taking as the general conclusion for that distance which we must at present rest in, that assigned in our article last cited, viz., 91,718,000 imperial (or 91,626,282 *geometrical*) miles, we find it equivalent to 23,222 polar semi-diameters of the earth, or ten million times that number of GEOMETRICAL CUBITS of twenty-five *geometrical inches* each.

(19.) Our next step is to the fixed stars, within whose sphere modern science has at length made good a footing, secure, though somewhat unsteady for the present. In conformity with the same principle of procedure, we here rest for our base of operations on our last and greatest *accessible* measured length, viz., the diameter of the earth's annual orbit, a base line of 183,000,000 miles, which, as the orbit is very nearly circular, presents itself (in some situation or other across it) perpendicularly to a line joining the sun and any selected star, so as to be seen unforeshortened from the star. As the earth at half-yearly intervals passes alternately from one to the other extremity of such a diameter, the visual line by